



Evolution of Supermassive Black Holes Binary

Introduction

The concept of gravitational waves is one of the most impressive predictions of Einstein's Theory of General Relativity. Gravitational waves are the space-time ripples propagating with the speed of light similarly to electromagnetic waves. Direct detection of gravitational waves is incredibly difficult, however, the first signal was detected on September 14, 2015, by LIGO and VIRGO collaborations.

Gravitational waves are emitted during the rapid motion of massive objects. The most powerful source of gravitational waves is the merging of two Supermassive Black Holes (SBH). Black holes predicted by Theory of General Relativity represent extremely compact objects which might have very large masses. Other specific properties of black holes will not be needed in the solution of this problem.

In the generally accepted theory of galaxies' evolution, it is supposed that there is SBH with the mass ranging from $10^5 - 10^9$ of Solar masses in the galaxy's center. Galaxies are huge stellar systems containing $10^{10} - 10^{11}$ stars. During their evolution, two galaxies can collide and merge into one. What happens to two SBHs initially located in their centers? The evolution of the SBH binary system can be divided into *three* main stages. At each stage SBHs approach each other, although the underlying physical phenomena differ. We will examine these phenomena separately in the first three parts of the problem. In the fourth part, we will use the obtained relations to calculate the total time of the SBH binary system evolution.

At the end of their evolution, two SBHs will eventually approach each other and merge into a single black hole. The merging process lasts about an hour and is accompanied by an intense burst of gravitational radiation. Future observatories like LISA will be able to detect this gravitational radiation. Still, the research on the SBH evolution is under way now, at the dawn of gravitational-wave astronomy.

General information

1. Express all your numerical answers in parsecs (pc) for the distances and giga-years (Gy) for the time intervals. We will use Solar mass (M_S) as a reference mass. You might need these values:

$$1 \text{ pc} = 3.1 \times 10^{16} \text{ m},$$

$$M_S = 2.0 \times 10^{30} \text{ kg},$$

$$t_H = 13.7 \text{ Gy}, \text{ age of the Universe},$$

$$G = 6.67 \times 10^{-11} \text{ N} \times \text{m}^2/\text{kg}^2,$$

$$c = 3.0 \times 10^8 \text{ m/s}.$$

2. When you encounter the word "**estimate**", you are not demanded the exact answer. It is sufficient to obtain a result that differs from the accurate one, not more than by a factor of 10. On the contrary, when you encounter the word "**find**", you are supposed to achieve the exact answer. The word "**calculate**" asks you to bring the numerical answer.

3. Throughout the problem, assume all stars in the galaxies to have the same mass that equals to Solar mass M_S .

4. Throughout the problem we will not take into account the effects of the Theory of General Relativity except gravitational waves emission. All stars and black holes are considered as point masses governed by Newton's gravitation law.



Part A. Dynamic Friction (1.6 points)

In this part we shall study the simplified model of the galaxy. You can ignore the velocities of the stars in the galaxy and assume the constant stellar concentration n . The characteristic size of the galaxy is R . The stellar concentration is small enough, so the stellar collisions are extremely rare and negligible. Let us consider a massive object with the mass $M \gg m$ moving with the velocity v through the galaxy. Surprisingly, the object experiences nonzero average force from the stars. This force slows the motion of the object and is called the force of dynamical friction for this reason. This part is devoted to the determination of this force.

- A.1** Let us work in the object's reference frame and consider the transit of one star with impact parameter b (fig. 1). Assume that 0.75pt

$$b \gg b_1 = \frac{GM}{v^2}. \quad (1)$$

The angular deflection of the star $\alpha = kb_1/b$, where k is some coefficient. Find the value of k . If you cannot find k , assume $k = 1$ hereafter.

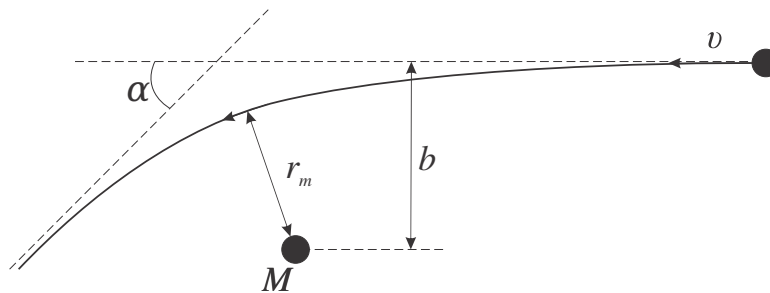


Fig. 1: The deflection of a star by object with mass M . The impact parameter is a , the minimal distance between the star and the object is r_m .

- A.2** Let Ox axis be directed along the object's velocity. Find the momentum component Δp_x transferred from the star to the object. 0.25pt

- A.3** Estimate the average force F_{DF} acting on the object by taking the average over impact parameter b . Neglect the contribution of the stars with impact parameters $b < b_1$. Assume the object to reside in the central part of a galaxy. Express F_{DF} in terms of M , v , R , G and stellar density $\rho = mn$. 0.4pt

- A.4** As you obtained in the previous task, the expression for F_{DF} includes the factor $\log R/b_1$, which we will denote further as $\log \Lambda$. Calculate the value of $\log \Lambda$ for $M = 10^8 M_\odot$, $R = 20 \text{ kpc} = 20 \times 10^3 \text{ pc}$ and velocity $v = 200 \text{ km/s}$. 0.2pt



Part B. Gravitational slingshot (3.0 points)

In this part, we will consider the system of two SBHs with equal masses $M \gg m$ located in the center of the galaxy. Let's call this system a **SBH binary**. We will assume that there are no stars near the SBH binary, each SBH has a circular orbit of radius a in the gravitational field of another SBH.

- B.1** Find the orbital velocity v_{bin} of each SBH. Find the total energy E of the SBH binary. Express it in terms of a , G and M . 0.25pt

There are a lot of stars at distances much larger than a from the binary. Stars travel along complex and diverse trajectories in the gravitational field of the whole galaxy. The motion of the stars can be considered chaotic, like the motion of the molecules in an ideal gas. Let us assume that stars' velocities have equal magnitudes $\sigma \ll v_{bin}$ and their average mass density is ρ . In this case dynamical friction is no longer affecting the SBH binary and energy losses are caused by other phenomenon.

- B.2** Let us solve a related problem. Let a star of mass m transit by a point mass $M_2 \gg m$ being at rest. The minimal distance between the star and the point during the transit is r_m . The velocity of the star at large distance is σ . Find the exact value of impact parameter b . 0.5pt

If a star approaches the SBH binary for a distance about a , it participates in a complex 3-body interaction with the binary that almost always results in a star being shot out with the velocity about v_{bin} (the velocity of the star at the large distance after interaction). We will call such a strong interaction a collision of a star with the SBH binary. Acceleration and the shot of the star after the collision is called "**gravitational slingshot**".

- B.3** Estimate the characteristic time Δt between two successive collisions of the SBH binary with stars. Take into account that $\sigma \ll v_{bin}$. 1.0pt

- B.4** Estimate the SBH binary energy loss rate dE/dt . Estimate the radius variation rate da/dt . Express it in terms of a , ρ , σ , G . 0.25pt

- B.5** Let us denote the initial radius of the system as a_1 . Estimate the time T_{SS} for the radius to decrease by a factor of 2 due to "gravitational slingshot". Calculate T_{SS} for $\sigma = 200$ km/s, $a_1 = 1$ pc, $\rho = 10^4 M_\odot/\text{pc}^3$. 1.0pt



Part C. Emission of gravitational waves (1.0 points)

In this part we shall study the SBH binary with equal masses which doesn't interact with the stars. Even in this case the system loses the energy due to gravitational waves emission. The energy loss rate due to gravitational waves is

$$\frac{dE}{dt} = -\frac{1024}{5} \frac{G}{c^5} (\omega^3 I)^2, \quad (2)$$

where ω is angular velocity of the binary, and $I = 2Ma^2$ is quadrupole moment of the system.

- C.1** Find the SBH binary radius variation rate da/dt due to the emission of gravitational waves. 0.2pt

When the orbit radius of the SBH binary a becomes close to the gravitational radius of the black hole:

$$r_g = \frac{2GM}{c^2}, \quad (3)$$

two SBHs quickly merge.

- C.2** Let us denote the initial radius of the system as $a_2 \gg r_g$. Estimate the time T_{GW} it takes for the SBH binary to shrink to the radius about r_g due to the emission of gravitational waves. Express T_{GW} as a function of a_2 , M , c and G . 0.7pt

- C.3** Calculate the initial radius a_H of the SBH binary with masses $M = 10^8 M_S$ if it takes it the age of the Universe to merge: $T_{GW} = t_H$. 0.1pt



Part D. Full evolution (4.4 points)

In this part we will use the results obtained above. Let us consider the **real** astrophysical situation. Two galaxies having SBH of mass $M = 10^8 M_S$ in their centers merged into a new stellar system. Let the new galaxy be spherically symmetrical with radius $R = 20 \text{ kpc} = 20 \times 10^3 \text{ pc}$. Let us assume that stellar density varies with radius r to the galaxy center as

$$\rho(r) = \frac{\sigma^2}{4\pi G r^2}, \quad (4)$$

where $\sigma = 200 \text{ km/s}$.

- D.1** Let the body move in circular orbit of radius $a < R$ in gravitational field of the stars. Neglect the force of dynamical friction and find the velocity v of the body. 0.25pt

Immediately after the merging of galaxies two SBHs have arbitrary positions inside the new galaxy and do not affect each other. Let's consider one SBH. We assume it moves in a circular orbit of radius $a < R$

- D.2** Estimate the orbit radius variation rate da/dt . In part A we ignored the velocities of the stars. Although stars are moving in the real galaxy, you can use the relation obtained in A.3 for the estimation. You should use the density $\rho(r)$ from equation (4). Assume $\log \Lambda$ to be a constant calculated in A.4. 0.75pt

After a certain time, two SBH will approach the center of the galaxy. Let two SBH move in a circular orbit of radius a around the center galaxy in the gravitational field of the stars.

- D.3** Estimate the critical radius a_1 at which gravitational interaction between two SBHs is no longer negligible and calculate it. We will say that at this moment two SBHs form a binary system (fig. 2). 0.3pt

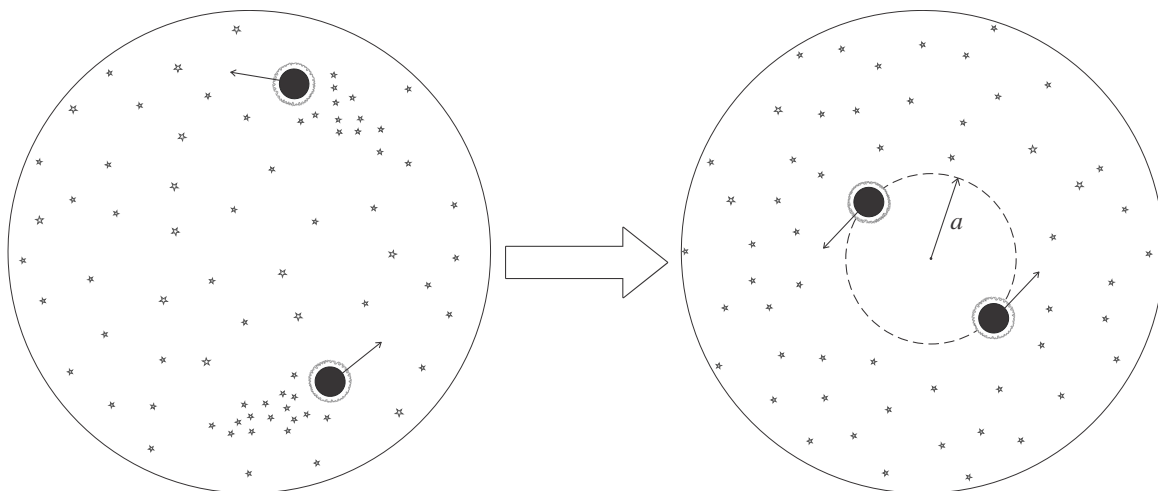


Fig. 2: The evolution of SBHs before and after the formation of the binary system



- D.4** Let us assume that after the merging of galaxies two SBHs were at distances $a_0 = 2 \text{ kpc} = 2 \times 10^3 \text{ pc}$ from the galaxy center. Calculate the time T_1 it takes for two SBH to form a binary due to dynamical friction. 0.75pt

After forming the binary, two SBHs shoot away all the stars from the center of the galaxy and stay there alone. Since this moment, dynamical friction becomes ineffective and the binary starts to lose the energy because of the slingshot effect. You can assume that the velocities of the stars around the binary are σ and the stellar density is $\rho_1 = \rho(a_1)$ from the equation (4). Slingshot effect shrinks the radius of the system drastically and after some time the system starts to lose energy mostly due to the radiation of gravitational waves.

- D.5** When the binary radius is less than some value $a < a_2$ the energy loss is caused by gravitational waves emission. Estimate the a_2 value and calculate it. 0.3pt

- D.6** Estimate the time T_2 of the binary radius reduction from a_1 to a_2 (the slingshot stage). Estimate the time T_3 of binary radius reduction from a_2 to almost zero (the stage of gravitational waves emission). 1.75pt

- D.7** For the parameters given above, calculate the total time T_{ev} of two SBH evolution from galaxies merging to SBH merging. 0.3pt

Conclusion. For a long time astrophysicists have been thinking the SBH binary evolution stops at the slingshot stage, since the binary has shot out all the stars with small impact parameters which might collide with it. It appeared that two SBH would never merge. This fact was called **the final parsec problem**.

Real galaxies have complicated asymmetric shapes. Few years ago it was found that in galaxies of complex shapes the stars with small impact parameters appear again and again. The SBH binary continues to lose energy, but slower than our estimation gives. The final parsec problem was successfully solved.